

CEvNS Signal





Coherent ν -Nucleus Scattering

47 years ago, Coherent Elastic Neutrino-Nucleus Scattering (CEvNS) was predicted with the realization of the neutral weak current.

D. Z. Freedman, PRD 9 (5) 1974

- Neutrino scatters coherently off all
 Nucleons → cross section enhancement:
 σ ∝ N²
- Initial and final states must be identical: Neutral Current elastic scattering
- Nucleons must recoil in phase →low momentum transfer qR <1 → *very* low energy nuclear recoil



Neutron number

CEvNS: A collaborative community



CEvNS: Exciting Decade ahead of us

| | 2017 | 2018 | 2019 | 2020 | 2021 | 2022 | 2023 | 2024 | 2025 | 2026 | 2027 | 2028 | 2029 | 2030's |
|-----------------------------------|------------------|-------------------|------------|-------------------|---------------|------------------|---------------|------------------|-----------|--------------|------|----------------|------|--------|
| Reactor Neutrino Sourc | ces | | | | | | | | | | | | | |
| Hartlepool Site | | | | BG studies | | | | | | | | | | |
| Angra Site | | | | location & shield | ding upgrades | | | | | | | | | |
| CHOOZ VNS Site | | BG studies | | Site Pre | paration | Operations | | | | | | | | |
| CONNIE | 40 g | | | 10 g skipper | 100g skipper | | 1 kg skipper | | | 10kg skipper | | | | |
| CONUS | | | | | | | | | | | | | | |
| MINER | | | | | | | | | | | | | | |
| RED-100 | | | | | | | | | | | | | | |
| TEXONO | | | PIRE R8 | kD | | | | | | | | | | |
| NUCLEUS | | | | | | 10g CaWO4 & Al20 | 3 | kg-scale1kg: Ge- | + Si | | | | | |
| XENON | | | | | | | | | | | | | | |
| NEWSG | | | Fea | asability Studies | | | | | | | | | | |
| RICOCHET | | | _ | | | | | | | | | | | |
| LAr | | | Det | tector R&D | | | | | | | | | | |
| Spallation Neutrino Sou | urces | | | | | | | | | | | | | |
| SNS FTS | 1.1 MW | 1.4 | MW upgrade | | | 1.7 MW upgrad | e | Ep=1.3 GeV Upg | rade | | | 2.0 MW upgrade | | |
| SNS STS | | | | | | | | | | | | | | |
| Lujan Center | | | | | | | 30 ns Upgrade | | | | | | | |
| ESS COLIFORNIT COL | | | | | | | | | | | | | | |
| COHERENT - CSI | 20 ko/d | terver of d | | | | 75 | | | | 10 | | | | |
| - LAr | 20 KeV 1 | Inreshold | | 10 40 | | /5 | окд | 100 1/2 | | 10 | | | | |
| - Ge | | | | тока | 2 2 T | | | 100 Kg | Cruce Mal | | | | | |
| | | | | | 5.5 1 | | 100 kg | | Cry0 Nai | | | | | |
| CCM - 1st detector | | | | | | | 100 Kg | | | | | | | |
| - 2nd detector | | | | | | | | | | | | | | |
| 8B Solar Neutrinos | 1 | 1 | 1 | 1 | | | | | | | | 1 | | |
| Xenon NT | | | | | | | | | | | | | | |
| LZ | | | | | | | | | | | | | | |
| SuperCDMS | | | | | | | | | | | | | | |
| Darkside-LM | | | | | | | | | | | | | | |
| Atmospheric and Diffus | se Supernova Neu | trino Backgrouino | d | | | | | | | | | | | |
| DARWIN | | | | | | | | | | | | | | |
| Galactic Supernova Neutrinos Only | | | | | | | | | | | | | | |
| Darkside-20k | | | | | | | | | | | | | | |
| ARGO | | | | | | | | | | | | | | |

Measuring the recoil: Quenching Factor



Improving the Precision for COHERENT CsI: New Quenching Factor

- Uncertainty on CEvNS rate
 improved from 25% to 4%
- Situation even more challenging for lower energy neutrinos at reactors



Another high energy example: COHERENT LAr

- Global analysis of world data
- Parametric model fit according to PDG prescription
- 2% uncertainty on CEvNS in ROI



Lowering thresholds can counteract QF uncertainty

Example: Germanium

- CONUS Rx Ge experiment producing data
- Growing confusion as to low-energy QF, with a major impact on expected rate above threshold

Is this the Migdal effect?





Example: Liquid Xenon

- LUX/LZ measurements at TUNL had dramatic impact on predicted Rx CEvNS signal
- Measurements continuing in LXe and LAr with CHILLAX



Energy (keV)

10



Example Silicon

- Dramatic drop below model prediction observed in silicon semiconductors. Implies significant impact
- Challenging measurement campaign to confirm with SuperCDMS Si detectors at TUNL with 55 keV ± 1 keV neutron beam



| ſ | Energy | $CE\nu NS$ -rate | $CE\nu NS$ -rate | 95% C. L. |
|---|---------------|------------------|------------------|-----------|
| | range (keV) | Lindhard | Chavarria | from data |
| ſ | 0.075 - 0.275 | 11.4 | 4.8 | 197 |
| | 0.275 - 0.475 | 3.6 | 1.3 | 109 |
| | 0.475 - 0.675 | 0.8 | 0.3 | 47 |

TABLE I. Expected rate from $CE\nu NS$, in events/day/kg/keV, assuming quenching factors from Lindhard [57] and Chavarria [52] together with the 95% CL limit from the data presented in this paper.





FIG. 20. CE ν NS event rate: 95% confidence level limit from the reactor on - off measurement (solid line) and neutrino signal expected from the Lindhard [57] (dotted line) and Chavarria [52] (dashed line) quenching factors.

Example: Nal[TI]

- Long history of Nal[TI] QF measurements
- Still a lot of variance
- Now testing variance across crystal manufacture technique (ANAIS and COSINE)
- Next up: varying dopant concentration & temperature





Example: Gaseous Ne

- Never before measured, so CEvNS signal unknown
- NEWS-G measurement at TUNL
- Challenge: scattering neutrons on gas





A Caveat: Bolometric Detectors

- Bolometric detectors (NUCLEUS, RICOCHET) do not suffer from QF loss of signal. Many of these also have extremely low thresholds
- Such detectors optimally suited to measure neutrinos
 <1.8 MeV IBD threshold
- Detector response should still be demonstrated



Actionable next steps

- More thorough measurement campaigns needed across all detector systems
- Unified approach to evaluating new and existing data
- Modeling remains a challenging (and perhaps unattainable) goal. Empirical measurements preferred
- Possible scoping study: should a decision be made on bolometric vs calorimetric detector systems? Do we avoid QF's all together? Such a study seems unlikely to be effective.

Trade-offs towards improved precision

- Small, purpose built detectors required to reduce multiple scattering.
- More beam-time needed at more facilities, but these facilities are not often run by agencies interested in this problem.
- Beam-time can be reduced with larger backing detector arrays, which also improve fidelity of results.
- But reduction of systematic uncertainties likely requires tuning of beam energy, as well as scanning over detector characteristics (e.g. doping concentration). Will require more beam time.
- Can move to higher neutron flux, but often this is also controls the neutron energies (e.g. 55 keV SuperCDMS run).
- Control of kinematics via neutron collimation, energy tuning, monochromaticity and beam pulsing will burn down systematic uncertainties.